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SCOPE Survey Description

SCOPE: SCUBA-2 Continuum Observations of Pre-protostellar Evolution – Survey Description and Compact Source Catalogue


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ABSTRACT

We present the first release of the data and compact-source catalogue for the JCMT Large Program SCUBA-2 Continuum Observations of Pre-protostellar Evolution (SCOPE). SCOPE consists of 850-μm continuum observations of 1235 Planck Galactic Cold Clumps (PGCCs) made with the Submillimetre Common-User Bolometer Array 2 on the James Clerk Maxwell Telescope. These data are at an angular resolution of 14.4 arcsec, significantly improving upon the 353-GHz resolution of Planck at 5 arcmin, and allowing for a catalogue of 3528 compact sources in 558 PGCCs. We find that the detected PGCCs have significant sub-structure, with 61 per cent of detected PGCCs having 3 or more compact sources, with filamentary structure also prevalent within the sample. A detection rate of 45 per cent is found across the survey, which is 95 per cent complete to Planck column densities of $N_{\text{H}_2} > 5 \times 10^{21}$ cm$^{-2}$. By positionally associating the SCOPE compact sources with YSOs, the star formation efficiency, as measured by the ratio of luminosity to mass, in nearby clouds is found to be similar to that in the more distant Galactic Plane, with the column density distributions also indistinguishable from each other.

Key words: surveys – stars: formation – ISM: clouds – submillimetre: ISM

1 INTRODUCTION

The Planck survey, with its primary goal of mapping the cosmic microwave background, covered the thermal emission from dust of $\sim 14$ K at wavelengths of 350 μm, 550 μm, and 850 μm. In the process of removing local, Galactic emission, a catalogue of 13188 Planck Galactic cold clumps (PGCCs) was compiled (Planck Collaboration et al. 2011a, 2016). By comparing the column densities and velocity widths of such clumps and those containing active star formation, the PGCCs were found to be significantly more quiescent and less evolved (Wu et al. 2012; Liu et al. 2013). The apparent quiescent nature of PGCCs makes them a valuable sample for studying the earliest stages of star formation, especially since they appear to have conditions suitable for star formation, with low dust temperatures of 6–20 K (Planck Collaboration et al. 2016). CO clumps have been detected towards PGCCs (e.g., Liu et al. 2013; Meng et al. 2013; Parikka et al. 2015; Zhang et al. 2016; Fehér et al. 2017), as well as detections of line emission from dense-gas tracers (Yuan et al. 2016). There can be, however, low levels of active star formation within PGCCs (Tröth et al. 2014; Liu et al. 2015; Tang et al. 2018; Yi et al. 2018; Zhang et al. 2018).

A large sample of prestellar cores and clumps needs to be studied to understand the evolution of cores and clumps after the formation of a young stellar object (YSO). In support of the effort to collect a large sample of pre- and protostellar cores and clumps, we present SCOPE, the SCUBA-2 Continuum Observations of Pre-protostellar Evolution Large Program. The project aims to test the earliest stages of star formation by observing 1235 PGCCs with the wide-field sub-mm bolometer camera, the Submillimetre Common-User Bolometer Array 2 (SCUBA-2; Holland et al. 2013) at the 850-μm wavelength beam size of 14.4 arcsec on the James Clerk Maxwell Telescope (JCMT). By observing with SCUBA-2 in the 850-μm continuum, matching the frequency of the Planck 353 GHz band, we can significantly improve over the 5″ resolution of the Planck observations of PGCCs. The 5″ resolution of Planck is prohibitive in that it hinders the positional cross matching of higher-resolution catalogues of YSOs (e.g. Wright et al. 2010; Gutermuth & Heyer 2015; Marton et al. 2016) and also doesn’t reveal the highly structured nature of the PGCCs (Juvela et al. 2012; Liu et al. 2012).

The resolution of SCUBA-2 at 850 μm resolves PGCCs at 1.5 kpc and 0.5 kpc to scales of 0.1 pc and 0.03 pc, respectively, the typical size of star-forming cores and clumps (e.g., Könyves et al. 2015), with 56 per cent and 43 per cent of sources falling within these distances, respectively (Planck Collaboration et al. 2016). It is important to resolve individual cores in the PGCCs, especially because of the close connection between the core mass spectrum and the stellar IMF (e.g., Simpson et al. 2008; Könyves et al. 2015; Montillaud et al. 2015).

The SCOPE survey was awarded 300 hours in the JCMT weather bands 3 and 4, which correspond to sky opacity values of $T_{225} = 0.08–0.2$. Observations began with three periods of pilot observations occurring in September 2014, and continued from December 2015 until July 2017. SCUBA-2 does observe the 450-μm band simultaneously with the 850-μm band but the weather bands available to SCOPE do not allow for reliable photometric calibration to be made on these shorter wavelength data.

1.1 SCOPE science goals

As previously mentioned, the Planck survey mapped the entire sky, therefore PGCCs cover all Galactic longitudes and latitudes, from the Galactic Plane to high-latitude clouds. Previous works suggest that the Galaxy’s spiral arms do not have much of an impact on star formation, other than collecting the source material together (Eden et al. 2012, 2013; Moore et al. 2012). Understanding star formation out of the Galactic Plane, however, is restricted to nearby clouds in the Gould’s Belt (e.g., Ward-Thompson et al. 2007; André et al. 2010). Some clouds at higher latitudes show signs of star formation (e.g., McGee 2008; Malinen et al. 2014; Kerp et al. 2016). Hence, high-latitude clouds could be contributing to Galactic star formation, both at present and in the future by providing source material for fresh star formation. Indeed, gas at high latitudes could be a part of a Galactic fountain, where gas and dust is expelled into the Galactic halo by supernovae and stellar winds (Bregman 1980). This material would then fall back to the Galactic Plane, cooling and condensing into atomic clouds, with observations detecting HI high-latitude clouds (Röhser et al. 2016b). These clouds then fall back onto the Plane,
SCOPE Survey Description

2 SCOPE OBSERVING STRATEGY

The SCOPE selection of PGCCs was chosen randomly from the full PGCC sample to sample varying Galactic environments (Liu et al. 2018a). Sources from the catalogue of PGCCs were excluded if they had already been observed at the JCMT with SCUBA-2 at a 6 mJy beam$^{-1}$ sensitivity. These observations formed part of the JCMT Legacy Surveys (Chrysostomou 2010), namely the JCMT Plane Survey (JPS; Moore et al. 2015; Eden et al. 2017 and the JCMT Gould Belt Survey (GBS; Ward-Thompson et al. 2007). The SCUBA-2 Ambitious Sky Survey (SASSy; MacKenzie et al. 2011; Nettek et al. 2017; Thompson et al., in preparation), covers half of the Galactic Plane but does not have the desired rms.

The PGCCs not previously observed at JCMT were placed into a 3-dimensional grid with each longitude bin 30 degrees wide, latitude bins of $|b|=0^\circ, 4^\circ, 10^\circ, 90^\circ$, and distance bins of 0 kpc, 0.2 kpc, 0.5 kpc, 1 kpc, 2 kpc, and 8 kpc. Sources with column densities $N_{\text{H}_2} < 2 \times 10^{21}$ cm$^{-2}$ chosen randomly within each bin, with all sources with column densities $N_{\text{H}_2} > 2 \times 10^{23}$ cm$^{-2}$ observed. This column density criterium is due to the pilot study of the initial 300 PGCCs, where the detection rate of PGCCs with $N_{\text{H}_2} < 10^{21}$ cm$^{-2}$ were dramatically lower. The JCMT has through SCOPE, and other archival data, covered nearly all PGCCs with column densities of $N_{\text{H}_2} > 2 \times 10^{21}$ cm$^{-2}$ within the declination range of $-30^\circ$ to $+35^\circ$, the most likely star-forming PGCCs (Wang et al., in preparation).

The properties of the observed sources are displayed in Fig. 1. The presented properties are Galactic longitude, Galactic latitude, temperature, distance, column density, mass, luminosity, major axis, and aspect ratio. The physical properties are not derived for all sources, and are as calculated by Planck Collaboration et al. (2016). These distributions are compared to the entire populations, as derived by Planck Collaboration et al. (2016). The biggest departure is with column density. We observed higher column density PGCCs than the average, which in turn skewed the SCOPE mass distribution to higher observed masses. Low column density ($N_{\text{H}_2} < 10^{21}$ cm$^{-2}$) PGCCs usually do not contain dense cores, based on our pilot observations, and thus are not interesting for star formation studies.

The average angular size of PGCCs is $8'$ (Planck Collaboration et al. 2016), therefore the CV Daisy mode of SCUBA-2 was used for these observations (Bintley et al. 2014). The CV Daisy mode is specifically designed for small, compact objects, with the telescope keeping to a circular pattern at $155''$ s$^{-1}$. Each SCOPE map takes 16 minutes to perform, with the CV Daisy mode producing 850-$\mu$m rms noise values of 6 mJy beam$^{-1}$ in the central $3'$, with 850-$\mu$m rms noise out to radii of $6'-10$ mJy beam$^{-1}$.

The SCOPE survey including pilot studies observed fields towards 1235 PGCCs in 1062 fields, with compact source detections towards $\sim 51$ per cent of fields, 45 per cent of PGCCs. The observed PGCCs in the SCOPE survey are displayed in their Galactic context in Fig. 2, with the 174 and 321 PGCCs observed by the JPS and GBS, respectively, overlaid. This paper will deal with only those observed as part of the SCOPE project.

The goal of SCOPE was to observe PGCCs in all Galactic environments. Comparison to the entire PGCC population and to sources in the GBS fields (Fig. 3) shows that SCOPE observations are representative of the high-column-density PGCC targets and complement the sources observed in GBS.

2.1 Complementary Observations

As well as the SCOPE survey, there are ongoing surveys at other facilities, giving complementary data to these PGCC detections (Liu et al. 2015, 2018a). The Taeduk Radio Astronomy Observatory (TRAO) Observations of Planck cold clumps (TOP) survey is observing $\sim 2000$ PGCCs in the rotational transition of $J = 1 - 0$ of CO isotopologues $^{12}$CO and $^{13}$CO at resolutions of 45-47 arcsec. A full description of this survey can be found in Liu et al. (2018a). Further $J = 1 - 0$ observations have been made at the Purple Mountain Observatory (PMO) at the same resolution. These observations will allow PGCCs to be put in the greater context of extended CO emission and structure.

The SMT (Submillimetre Telescope) “All-sky” Mapping of Planck Interstellar Nebulae in the Galaxy (SAMPLING;
Wang et al. (2018) survey follows up PGCCs in the \( J = 2 - 1 \) transition of \(^{12}\)CO and \(^{13}\)CO. The detection of two transitions will allow more accurate column densities to be calculated.

Further PI observations have occurred at the Nobeyama Radio Observatory 45-metre telescope, the 21-metre telescopes in Korean VLBI Network (KVN) and the Effelsberg 100-metre telescope in the dense-gas tracers like HCO\(^+\), N\(_2\)H\(^+\), HC\(_3\)N, CCS, DNC, HN\(^{13}\)C, N\(_2\)D\(^+\) and NH\(_3\) towards samples of SCUBA-2 dense cores in PGCCs. By observing these species, temperatures, depletion and deuteration fractions can be determined (Tatematsu et al. 2017).

### 3 DATA REDUCTION AND SCOPE DATA

Liu et al. (2018a) provides a full description of different data reduction methods employed and tested within the SCOPE survey.

The data reduction employed within the first SCOPE data release use the Dynamic Iterative Map-Maker (Chapin et al. 2013), part of the Starlink SMURF package (Jenness et al. 2011). These data make use of a 200\('\) spatial filter, with no use of external masking. This initial data reduction is then filtered for regions with a signal-to-noise ratio (SNR) less than 3, with these high-SNR regions used then as a mask for a further reduction. A full description of the masking process can be found in Mairs et al. (2015).

A flux conversion factor of 554 Jy beam\(^{-1}\) pW\(^{-1}\) is used to convert from the native units of pW to Jy beam\(^{-1}\). This value is \( \sim 3 \) per cent higher than the 537 Jy beam\(^{-1}\) pW\(^{-1}\) flux conversion factor recommended by Dempsey et al. (2013b), reflecting the pixel size (4\('\)) and data reduction method used by the SCOPE survey as the pixel size is a factor in the flux conversion factor equation.

Examples of two of the observed PGCCs are shown in Fig. 4, a complex, filamentary source, and a high-latitude cloud with simple morphology.

The mean rms within the central 12\('\) of the 1062 fields (this is the number of fields observed, containing the observed 1235 PGCCs) is 0.185 Jy arcsec\(^{-2}\), which corresponds to 43.9 mJy beam\(^{-1}\). This sensitivity is a factor of \( \sim 1.5 \) worse than the 25-31 mJy beam\(^{-1}\) rms of the JPS (Eden et al. 2017). Within
the central $3'$, the rms is found to be 0.028 Jy arcsec$^{-2}$, which corresponds to 6.65 mJy beam$^{-1}$.

4 COMPACT SOURCE CATALOGUE

4.1 Compact Source Extraction

Compact sources were extracted from the images using the FELLWALKER algorithm (FW; Berry 2015). Justification for the choice of source extraction algorithm and the parameters used are outlined in Moore et al. (2015) and Eden et al. (2017). The parameter FELLWALKER:MINPIX is adjusted to account for the larger pixel size used in the SCOPE reduction.

The FW algorithm is run on the SNR maps, with the mask produced by CUPID:FINDCLUMPS used to extract flux from the emission maps.

Compact sources were initially identified in 821 of the 1062 observed fields. Sources which had a peak SNR $< 5$ were rejected, as well as sources with an aspect ratio $> 5$. The sensitivity to extended, filamentary objects in the SCOPE survey will be explored in a future study (Fich et al., in preparation). These quality control cuts resulted in 3528 sources, with examples of the source extraction results for two fields shown in Fig. 4.

Table 1 contains a portion of the full SCOPE compact source catalogue. The columns are as follows: (1) SCOPE catalogue source name; (2) SCOPE Region; (3) and (4) Right Ascension and Declination (J2000) of the peak flux position within the SCOPE source; (5) and (6) Right Ascension and Declination (J2000) of the central point; (7–9) semi-major axis, semi-minor axis, and position angle, measured anticlockwise from Equatorial north, of the ellipse fit to the shape of the SCOPE source which are not deconvolved sizes; (10) effective radius of source, calculated by $\sqrt{\langle A^2 \rangle / \pi}$, where $A$ is the area of the source above the detection threshold of 3$\sigma_{\text{rms}}$; (11–12) peak flux density, in units of Jy beam$^{-1}$, and associated uncertainty; (13–14) integrated 850-µm flux, in units of Jy, and associated uncertainty and (15) signal-to-noise ratio (SNR) of the source, calculated from the peak flux density and the $\sigma_{\text{rms}}$ from the observed field. The uncertainties take account for errors in calibration, taken to be 5 per cent (Dempsey et al. 2013b), and uncertainties in the FCF value used, also taken to be 5 per cent. A full version is included in the Supporting Information. The 3528 sources were distributed across 558 PGCCs.

The FW routine was tested extensively within the JPS survey, with a 95 per cent recovery fraction of artificial sources found to be at approximately 50. A full explanation can be found in Eden et al. (2017). Therefore, we are assuming that this completeness limit is valid for the observed SCOPE sources within the Galactic Plane. An improved recovery fraction of 99 per cent was found in two SCOPE fields at $|b| > 30\degree$.

4.2 Recovered flux densities

The FW algorithm is well understood (Eden et al. 2017). Therefore, any differences in the recovered flux densities compared to other surveys are likely due to calibration issues. To test the recovered fluxes, the SCOPE fluxes were compared to those of the JPS survey (Eden et al. 2017), both positionally matched and the survey as a whole.

The distributions of peak intensities and integrated fluxes for both surveys are displayed in Fig. 5. The JPS survey source intensities have been converted into mJy arcsec$^{-2}$ and mJy in the peak intensity and integrated flux distributions, respectively. The peak intensity distribution of the JPS goes $\sim 2 \times$ deeper than SCOPE, which corresponds to the rms values of the respective surveys. The CV Daisy mode actually produces deeper observations in the central regions, with SCOPE having greater sensitivity in those regions. The peaks of the integrated flux distributions, however, are consistent with each other. By assuming single power-law laws for the tails of the distributions of the form $\Delta N/\Delta S \propto S^{-\alpha}$, values of $\alpha$ for the two distributions were found to be $\alpha = 2.10 \pm 0.13$ and $\alpha = 1.97 \pm 0.10$ for the peak intensity and integrated flux distributions, respectively, above limits of 0.5 mJy arcsec$^{-2}$ and 2 mJy for the peak intensity and integrated flux distributions, respectively. The SCOPE peak intensity distribution is consistent with that of the JPS ($\alpha = 2.24 \pm 0.12$) but the SCOPE integrated flux distribution is flatter than that of the JPS survey ($\alpha = 2.56 \pm 0.18$).

There is actually a slight overlap between the SCOPE and the JPS surveys. By positionally matching the two surveys within a JCMT beam of 14.4 arcsec, we find 83 matches. A comparison of peak intensities and integrated fluxes for sources in common are shown in Fig. 6. The two distributions show a slight discrepancy, with the integrated fluxes departing by a greater amount. The difference between the peak fluxes can be accounted for by the difference in pixel sizes (3 arcsec in the JPS, 4 arcsec in SCOPE). As reported in Mairs et al. (2015); Rumble et al. (2015), changing the pixel size, especially to smaller sizes, can change the peak value, with 3-arcsec pixels giving the most accurate peak fluxes. The difference in integrated fluxes is accounted for by larger sources in the SCOPE survey, with sources in the SCOPE survey found to have a mean size 1.32 times that of the reported JPS source size. A linear best fit to the relationship gives a gradient of 1.06 $\pm$ 0.05. The different FCFs will also account for some of the difference. However, the major difference comes from the sensitivity in the outer edges of the SCOPE maps, where these 83 sources are found. In these regions, the JPS is $\sim 2 \times$ more sensitive, thus causing the sources to be broken up in the JPS. A full explanation of this effect is contained in Eden et al. (2017) in the context of JPS and ATLASGAL (Schuller et al. 2009) comparisons.

4.3 Angular size distribution

The angular size distribution of the SCOPE sources is shown in Fig. 7. The plotted quantity is the major axis of the elliptical fit to the source, provided by FW. The reported sizes are not deconvolved sizes. The peak of the distribution, found at 35 arcsec, is in marked contrast to the peak at 8 arcmin found in the Planck catalogue (Planck Collaboration et al. 2016). This difference further exemplifies the inner substructure identified by the higher resolution SCOPE survey and the presence of multiple SCUBA-2 sources inside a single PGCC, which is highlighted in Fig. 8, where 61 per cent of detected PGCCs have 3 or more SCOPE sources.

We present the aspect ratios of the SCOPE sources in Fig. 9. We have overlaid the aspect ratios of the observed Planck PGCCs in this survey. Planck Collaboration et al. (2016) found that 40 per cent of sources had an aspect ratio of between 2 and 3, with a shoulder at those values not present in the SCOPE compact source catalogue. This shoulder in the PGCC catalogue aspect ratios is a hint at the substructure observed by SCOPE, with a higher aspect ratio pointing towards more filamentary structures. It is also a reflection of the nature of the sources extracted by the FW algorithm, which is attuned more towards extracting compact objects.

FW is part of the Starlink CUPID package outlined in Berry et al. (2007).
Figure 2. The distribution of observed PGCCs by the JCMT. The black points represent SCOPE sources, whilst the white and gold points were observed as part of the JPS and GBS, respectively. The underlying image is the Planck 353-GHz (850-µm) intensity map.

Figure 3. Histograms of column density of PGCCs with the entire population represented by the blue histogram, the grey histogram representing those PGCCs observed by SCOPE and the green hashed histogram covering those PGCCs observed by the JCMT in the Gould’s Belt Survey (Ward-Thompson et al. 2007).

5 SCOPE SURVEY AND DATA ACCESS

The SCOPE data products can be downloaded from the Canadian Astronomy Data Centre’s JCMT Science Archive\(^2\), with the proposal IDs MJLSY14B, M15AI05, M15BI06, and M16AL003. These IDs correspond to observations taken in the SASSy survey, two PI proposals, all of which formed the pilot observations, and the SCOPE survey, respectively. As well as these data, the raw observation data can also be downloaded from the same location.

The full source compact source catalogue is available as Supporting Information.

6 RESULTS

6.1 Detection statistics

The PGCCs observed in the SCOPE survey were chosen to sample across the spectrum of a host of statistics, as displayed in Fig.1, with a bias towards the highest column density sources, as described earlier. The detection rate of each statistic is shown in Fig.10. Six high-latitude PGCCs with detections (|b| > 30\(^\circ\)) were

Figure 4. Examples of observed PGCCs in the SCOPE survey. Left panel: A complicated, filamentary source at Galactic coordinates $\ell = 36\degree\,62, b = -0\degree\,11$, PGCC_G36.62-0.11. Right panel: A high-latitude PGCC, positioned at $\ell = 6\degree\,04, b = 36\degree\,77$, PGCC_G6.04+36.77 (Liu et al., in preparation). The intensity scale in each image is $\text{mJy arcsec}^{-2}$ and the white ellipses represent the elliptical fits to the FELLWALKER extractions within the observed field.

Figure 5. Peak and integrated flux distributions for the SCOPE survey (grey, filled histogram) compared to the JPS (blue histogram) in the left and right panels, respectively. The least-squares fit to the SCOPE distributions are indicated by the red, dashed line.

excluded from the statistics as they were found to be associated with lensed galaxies (Liu et al., in preparation).

The longitude detection rates are approximately equal to the overall detection rate (0.46), with lower rates found in the central 20 degrees ($|b| < 2\degree$). In these longitudes, higher latitude sources are observed, with lower detection rates found outside of the central 4 degrees ($|b| > 2\degree$) of the Galactic Plane. In the latitude range of $|b| < 2\degree$, the detection rate is slightly higher than the overall rate, approximately 0.50. This is also reflected in the detection rate as a function of distance. The detection rate at higher distances, those classically taken to be in the Galactic Plane, are well detected. However, the local, higher latitude sources are not detected at as higher rate.

The column densities are 95 per cent complete above densities of $N_{\text{H}_2} > 5 \times 10^{21} \text{ cm}^{-2}$. This detection rate corresponds to the threshold for star formation found in studies of nearby star-forming clouds (André et al. 2010; Heiderman et al. 2010; Lada et al. 2010; André et al. 2014) Some bins are subject to low detection rates due to low number statistics. The mass and luminosities are complete to a 95 per cent rate above $5 \times 10^3 \text{ M}_\odot$ and $1 \times 10^3 \text{ L}_\odot$, respectively.

The column density detection rate does not account for the selection bias imposed on the initial sample. When comparing the entire PGCC population, the observed SCOPE sample, and the detected PGCCs in SCOPE, we can see that the highest column density sources are confined to the lower latitudes, which are the
PGCCs with the largest distances, as seen in Fig. 11. Deeper observations will be required to trace the lowest column density objects.

The major axes and aspect ratios fluctuate around the actual detection rate of the SCOPE survey.

The detection rate of the temperatures is skewed by the lack of derived temperatures by Planck Collaboration et al. (2016). 619 of the 1235 observed PGCCs do not have derived temperatures, with 360 of those 619 detected in SCOPE. The detection rate at the lower temperatures of less than 15 K is 0.15.

6.2 Star formation out of the Galactic Plane

The SCOPE survey gives a sample of potentially star-forming cores and clumps in different Galactic environments. The YSO catalogue of Marton et al. (2016) provides an all-sky catalogue derived from the AllWISE catalogue (Cutri et al. 2013). By positionally matching these two catalogues, we find 865 YSOs located within the map extents of 201 observed PGCCs that have detected compact sources.

The YSO catalogue of Marton et al. (2016) contains the magnitudes of the 4 WISE bands (3.4 µm, 4.6 µm, 12 µm, and 22 µm) as well as the $J$, $H$, and $K$ bands from the positional matching of 2MASS Point Sources (Cutri et al. 2003). These 7 bands can be used to calculate the luminosities of each YSO, and therefore the total YSO luminosity associated with a PGCC. The luminosities are calculated, once the magnitudes are converted into fluxes, using a trapezium rule estimation in log-log space, which was shown to provide a good approximation of the luminosity (Eden et al. 2015), with other studies using this method (e.g., Veneziani et al. 2013).

These luminosities, $L$, can be compared to the masses, $M$, of the SCOPE detected YSOs to determine the ratio of $L/M$, a measure of the current star formation and an indicator of the evolutionary state of that star formation (e.g., Elia et al. 2017; Urquhart et al. 2018). Comparing this ratio between the nearby sources, complementary to the Gould’s Belt, and the more distant Galactic Plane...
can determine whether the star formation in either environment is at a different evolutionary stage (on average).

The masses were derived from the total emission within the SCOPE maps. The masses in the SCOPE maps are estimated using the optically thin approximation:

\[ M = \frac{S_v D^2}{\kappa_v B_v(T_d)} \]

where \( S_v \) is the integrated flux of the emission, \( D \) is the distance (omitted in the calculation for the ratio of \( L/M \)), \( \kappa_v \) is the mass absorption coefficient taken to be 0.01 cm\(^2\) g\(^{-1}\) (Mitchell et al. 2001) which accounts for a gas-to-dust ratio of 100, and \( B_v(T_d) \) is the Planck function evaluated at temperature, \( T_d \), taken to be 13 K, the peak of the SCOPE-observed PGCC-temperature distribution.

The sample of 201 PGCCs with a YSO was split into two populations, a Gould’s Belt-like population and the distant Galactic Plane. This was done by using a distance cut of 0.5 kpc, as this is taken to be the furthest Gould’s Belt sources (Ward-Thompson et al. 2007), with distances below this in the Gould’s Belt, and greater distances in the Galactic Plane. From hereafter, these populations will be referred to as the nearby and distant samples, respectively. The distances were estimated in three ways. The first was to take the distances as derived by Dame et al. (2001). This accounted for 91 PGCCs. The second was to extract spectra from spectral line data from existing surveys (Dame et al. 2001; Jackson et al. 2006; Dempsey et al. 2013a; Rigby et al. 2016), with sources in the Plane compared to the Galactic rotation curve of Brand & Blitz (1993). This accounted for a further 108 PGCCs. The final two were well out of the Plane at high latitudes, and were assumed to be local.

We found 114 and 87 PGCCs in the distant and nearby samples, respectively. The distribution of the \( L/M \) ratios for each of these populations is shown in Fig. 12. The mean values of the two populations are 0.78 \( \pm 0.16 \) L\(_{\odot}\)/M\(_{\odot}\) and 1.14 \( \pm 0.30 \) L\(_{\odot}\)/M\(_{\odot}\) for the distant and nearby samples, respectively, with median values of 0.21 \( \pm 0.16 \) L\(_{\odot}\)/M\(_{\odot}\) and 0.32 \( \pm 0.28 \) L\(_{\odot}\)/M\(_{\odot}\), respectively. The populations are consistent with each other, within the errors, indicating that the star formation in the distant Galactic Plane is at roughly the same evolutionary stage as that within the nearby Gould’s Belt-like sample. A Kolomogorov–Smirnov (K–S) test of the two subsamples finds a 25 per cent chance they are not drawn from the same distribution. We therefore cannot strongly reject the null hypothesis that they are not drawn from the same sample and that the distributions at different distances are different. The sample of the Galactic Plane PGCCs will be extended when JPS sources are also included, and this analysis will be the subject of a further study (Eden et al., in preparation).

6.3 Column densities of SCOPE sources

The column densities of the SCOPE sources in the nearby environments and the distant Galactic Plane are compared. The column densities were calculated using the following:

\[ N_H = \frac{S_v \text{peak}}{B_v(T_d) \Omega_\theta \kappa_v \nu \mu} \]

where \( S_v \text{peak} \) is the peak intensity, \( B_v(T_d) \) and \( \kappa_v \) are as defined above, \( \Omega_\theta \) is the solid angle of the beam, \( \mu \) is the mass of a hydrogen atom, and \( \nu \) is the mean mass per hydrogen molecule, taken to be 2.8 (Kauffmann et al. 2008).

The distribution of column densities for the entire SCOPE sample, as well as the two subsamples are displayed in Fig. 13 (top, middle), with the cumulative distribution also included. A K–S test of the two subsamples gives a \( \sim 2.5 \sigma \) or a 2.5 per cent chance that they are not drawn from the same population. We therefore cannot strongly reject the null hypothesis that they are not drawn from the same sample.

When comparing only the star-forming samples, as shown in Fig. 13 (bottom), the K–S test gives a 94 per cent result that the nearby and distant Galactic Plane SCOPE sources are drawn from the same population. We can assume, however, that the total star-forming sample is considerably different from the entire SCOPE sample due to a K–S test giving a probability of their being the same of 0.001. This result is consistent with that of Urquhart et al. (2014), who found that the star-forming clumps in the ATLASGAL survey had a considerably different column-density distribution than that of the entire population.

7 SUMMARY

We present the first data release of the SCOPE survey, presenting the data images and a compact source catalogue. The data consist of observations of 1235 Planck Galactic Cold Clumps (PGCCs) at angular resolutions of 14.4 arcsec, significantly improving upon the 5-arcmin resolution of Planck. The improved resolution reveals significant substructure within these sources, reflected by a compact source catalogue consisting of 3528 sources. The data are downloadable from the CADC, with the compact source catalogue included as Supporting Information to this article.

The compact source catalogue was produced using the FELLWALKER algorithm, reflecting the same method used within the JCMT Plane Survey (JPS; Eden et al. 2017). A comparison of peak intensities and integrated fluxes in overlapping sources between the JPS and SCOPE surveys shows slight discrepancies, but otherwise good agreement. The SCOPE sources are significantly smaller than those within the Planck catalogue, with peaks of the angular-size distribution found at 35 arcsec compared to 8 arcmin.

The detection rate of PGCCs within the SCOPE survey is 45 per cent, with 558 PGCCs detected of the 1235 observed. The survey is 95 complete in PGCCs with column-den-
... comment, they have greatly improved the quality of the paper.

Figure 10. Distributions of the observed, Planck-derived statistics of SCOPE sources (grey histograms) overlaid with the detected PGCCs (yellow hashed histogram). The ratio of the two histograms is overlaid by the blue dotted line. The black dashed line represents the overall detection rate. As in Fig. 1, the top row displays the Galactic longitude, Galactic latitude, and temperature, in the left, central, and right panels, respectively. The middle row are the distance, column density, and mass, with the bottom row showing the luminosity, major axis, and aspect ratio.

stellar \( N_{H_2} > 5 \times 10^{21} \text{ cm}^{-2} \), and to masses and luminosities of \( 5 \times 10^3 \text{ M}_\odot \), and \( 1 \times 10^3 \text{ L}_\odot \), respectively.

By positionally matching the SCOPE compact sources with YSOs from the WISE catalogue, and splitting the sample into sources that are within 0.5 kpc and those at greater distances, we found that the ratio of \( L/M \) is consistent between these samples. The column densities of these two samples of SCOPE sources are also consistent. The distribution of column densities of star-forming sources, however, were found to be significantly different from those of the whole SCOPE compact source catalogue.

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Figure 11. Latitude of PGCCs against column density (left panel) and distance (right panel). The blue circles are the entire PGCC sample from Planck Collaboration et al. (2016), the black plus symbols are the observed SCOPE PGCCs, with the yellow squares representing the detected SCOPE PGCCs.

Figure 12. Histogram of the $L/M$ ratios in SCOPE detected PGCCs in the further (grey) and in the nearby (green) Galactic Plane.

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Figure 13. Top panel: column density distribution of all SCOPE compact sources (grey histogram), with those in the distant Galactic Plane ($d > 500$ pc; red histogram) and nearby sample ($d < 500$ pc; green histogram). Middle panel: the cumulative distributions of the same histograms with the dashed black line, dotted red line, and green solid line representing, the whole sample, in the distant Galactic Plane and nearby Plane, respectively. Bottom panel: the cumulative distributions of the star-forming samples, with the dashed black line representing the total sample (as in the middle panel, includes non-star-forming sources), and the dotted red lines and green solid lines represent the star-forming sources in distant and nearby environments, respectively.
Table 1. The SCOPE compact source catalogue. The columns are as defined in Section 4.1.

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<th>Name</th>
<th>Region</th>
<th>RA&lt;sub&gt;peak&lt;/sub&gt; (°)</th>
<th>Dec&lt;sub&gt;peak&lt;/sub&gt; (°)</th>
<th>RA&lt;sub&gt;cen&lt;/sub&gt; (°)</th>
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<th>σ&lt;sub&gt;min&lt;/sub&gt; (°)</th>
<th>PA (°)</th>
<th>S&lt;sub&gt;peak&lt;/sub&gt; (Jy arcsec&lt;sup&gt;-2&lt;/sup&gt;)</th>
<th>ΔS&lt;sub&gt;peak&lt;/sub&gt; (Jy arcsec&lt;sup&gt;-2&lt;/sup&gt;)</th>
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Note: Only a small portion of the catalogue is shown here. The entire catalogue is available in the Supporting Information.
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