Herschel: the first science highlights

LETTER TO THE EDITOR

Herschel-SPIRE spectroscopy of the DR21 molecular cloud core*


(Affiliations are available in the online edition)

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ABSTRACT

We present far-infrared spectra and maps of the DR21 molecular cloud core between 196 and 671 μm, using the Herschel-SPIRE spectrometer. Nineteen molecular lines originating from CO, 13CO, HCO+ and H2O, plus lines of [N II] and [CI] were recorded, including several transitions not previously detected. The CO lines are excited in warm gas with Tkin ∼ 125 K and nH2 ∼ 7 × 104 cm−3, CO column density N(CO) ∼ 3.5 × 1018 cm−2 and a filling factor of ∼12%, and appear to trace gas associated with an outflow. The rotational temperature analysis incorporating observations from ground-based telescopes reveals an additional lower excitation CO component which has a temperature ∼78 K and N(CO) ∼ 4.5×1017 cm−2.

Key words. ISM: lines and bands – infrared: ISM – ISM: molecules

1. Introduction

We report observations of the far-IR spectrum of the DR21 molecular cloud core obtained with the Herschel satellite between 196 and 671 μm. The DR21 HII-region/molecular cloud is part of the Cygnus X complex of molecular clouds located at a distance of 1.7 kpc (Schneider et al. 2006). This region has been subject to numerous studies at different wavelengths (Richardson et al. 1988, Wilson & Mauersberger 1990; Liechti & Walmsley 1997; Schneider et al. 2006, 2010; Jakob et al. 2007). The main DR21 cloud core has a mass of ∼20000 M⊙ (Richardson et al. 1989), and contains one of the most energetic star formation outflows detected, with an outflow mass of ∼3000 M⊙ (Garden et al. 1991; Cruz-González et al. 2010).

2. SPIRE observations

2.1. Spectra

We present science demonstration phase (SDP) observations obtained with ESA’s Herschel Space Observatory (Pilbratt et al. 2010), using the Spectral and Photometric Imaging Receiver (SPIRE – Griffin et al. 2010). The calibration and characteristics of SPIRE have been described by Swinyard et al. (2010). SPIRE was operated as an imaging Fourier-transform spectrometer (FTS) in the high resolution mode (∆λ/λ = 1000 (≈300 km s−1 at 250 μm) sampling across an approximately circular field of view with an unvignetted diameter of 2.6′. This means that the line profiles are unresolved.

The sky footprint is formed by two detector arrays: the 19 pixel SLW array (671–303 μm) and the 37 pixel SSW array (313–194 μm), with beam widths varying from 17″ at 194 μm to 42″ at 671 μm, with uncertainties of ±7–10% (Griffin et al. 2010). The integration time was 1065 s, summed from two separate observations. The current best estimates of the absolute uncertainties for the FTS detectors are 10–20% for the SSW detectors, and ∼30% for the SLW detectors (Swinyard et al. 2010).

The unapodised FTS spectra provide the highest spectral resolution, with a classical instrumental sinc function line shape. A spectral line fitting routine was developed for extracting line parameters (Jones et al. 2009). This fits a continuum (either a low order polynomial or a blackbody variant) using the Levenberg-Marquardt least squares method. The fitting procedure weights the spectral intensity at a given frequency of an averaged spectrum by the statistical uncertainty at that frequency, returning line centers, intensities, line widths and their associated fit errors.

2.2. Maps

The SPIRE observations sparsely sample the field of view, although there are calibration uncertainties for the outer ring of detectors at the edges of both arrays that are not yet fully characterised. To provide a first look at the relative distributions in the various species, we have interpolated the fluxes of individual
The maps of selected species are shown in Fig. 2. The CO lines in both detector arrays show a prominent central peak, with extensions to the east and west along the well-known outflow. This has been assumed to be associated with outflowing gas with \( T_{\text{ex}} \sim 2000 \) K and \( H(\text{H}_2) \sim 1 \times 10^{19} \) cm\(^{-2} \) from Garden et al. 1991). However, as will be seen in the high resolution JCMT observations (Fig. 3), the emission traced in the SPIRE maps is also clearly visible in the relatively low excitation CO \( J = 3-2 \) data, suggesting that there may be a mixture of low and high excitation gas present. This is confirmed in Fig. 3, where similar extensions of the ambient gas are present in the JCMT CO \( J = 3-2 \) map, and that of Schneider et al. (2010).

This is not unexpected, as this outflow appears to have a very large mass of several thousand \( M_\odot \), and presumably the high velocity gas phase overlaps (or may co-exist with) ambient material. The SPIRE maps also show that the \( ^3\text{P}_1-^3\text{P}_0 \) atomic carbon line has a similar spatial distribution to that of CO. By contrast, the \( \text{H}_2\text{O} \) and [N II] lines appear to be more compact and centered close to the DR21 cloud core, although the [N II] distribution is elongated to the east – observations with higher signal to noise and better sampling are needed for more detailed comparison.

### 2.3. JCMT CO \( J = 3-2 \) observations

CO \( J = 3-2 \) JCMT archival data (programme M07AU01) with a 15″ beam and spectral of 0.05 km s\(^{-1} \) are shown in Fig. 3, from a 4.5 h integration using the HARP array receiver. The area covered by the SPIRE footprint (Fig. 2) is shown as a white square. The JCMT observations clearly trace the outflow which runs from the NE-SW from DR21 from the centre of the white box. The JCMT map also reveals a prominent north-south ridge that includes CO peaks associated with the well-studied sources DR21(OH) and DR21-FIR1.

The [N II] distribution is coincident with the distribution of high velocity gas and shocked \( \text{H}_2 \) (Garden et al. 1991).
3. Modeling the CO lines

The most extensive modeling of the CO toward DR21 is by Richardson et al. (1986, 1988), Wilson et al. (1990), Schneider et al. (2006) and Jakob et al. (2007). Richardson et al. (1986, 1988) presented a multiphase model with gas densities spanning the range $10^2$–$10^6$ cm$^{-3}$, and gas temperatures in the low temperature component $\geq$30 K. Jakob et al. (2007) confirmed this using KOSMA and ISO observations, finding and additional warm phase component with $T_{\text{kin}} \sim 80$–150 K and clump density $n_{\text{H}_2} \sim 10^4$–$10^6$ cm$^{-3}$.

We initially constructed a rotational temperature diagram for the SPIRE CO and $^{13}$CO lines. These were augmented with the JCMT CO line from Sect. 2.3, plus IRAM CO $J = 2$–1 observations (Schneider et al. 2010), with suitable beam size corrections. The rotational temperature diagram is shown in Fig. 4.

Both species show evidence for two gas components, a lower temperature phase with a rotational temperature $T_{\text{rot}} = 78$ K and total CO column density $N$(CO) $\sim 4.5 \times 10^{21}$ cm$^{-2}$, in addition to a higher temperature component with $T_{\text{rot}} = 185$ K and $N$(CO) $\sim 9.7 \times 10^{17}$ cm$^{-2}$. The $^{13}$CO lines are more limited and noisy, with the SPIRE lines indicating an intermediate temperature phase having $T_{\text{rot}} = 109$ K and $N$(13CO) $\sim 8.8 \times 10^{16}$ cm$^{-2}$. The data for $^{13}$CO also show evidence for a low temperature component, although this relies on comparison with low frequency ground based data (JCMT, IRAM) obtained with different beam sizes. Such a result is expected, since the observations probe deeper into the PDR of each clump in $^{13}$CO than in CO.

There are several problems with the rotational temperature approach, including wavelength dependent beam size corrections, opacity and calibration errors. These uncertainties can however be mitigated by i) taking ratios of the various CO line intensities on a single detector and using these to constrain the excitation conditions though our LVG modeling, and ii) using observations from the central pixel where the SSW and SLW beams are coincident and the calibration is well determined. This approach particularly mitigates against the beam size and calibration errors, since only flux ratios are being used to estimate the excitation conditions.

The model fit was made to the CO and $^{13}$CO lines using the off-line version of the RADEX LVG code (Van der Tak et al. 2007). The line ratios observed on the same detectors (hence the beam sizes are similar) were used to restrict the likely
excitation conditions. It proved difficult to find an unique single temperature model that simultaneously predicted the relative intensities of both isotopolues. However, the SPIRE data can be approximately reproduced by a single phase moderate temperature gas with $T_{\text{kin}} \sim 125$ K, volume density $\sim 7 \times 10^2$ cm$^{-3}$, with $N(\text{CO}) \sim 3.5 \times 10^{10}$ cm$^{-2}$, filling factor $\sim 12\%$, and a $[\text{C}]/[\text{CO}]$ ratio of 65. This model does however slightly overpredict the low-$J$ $^{13}\text{CO}$ 4–3 to 6–5 line intensities, compared to the $J = 7$–6 line. Changing the temperature and density from these conditions considerably worsened the high-$J$ CO line fits, although a more complex multiphase model, with appropriate (and uncertain) beam size corrections would improve the fit of the low $J$-lines. We have not attempted to fit to a PDR-model, as the data and calibration quality need to be improved if tests between models are to be made, and that this is beyond this first look paper.

4. Modeling the H$_2$O and [N II] lines

An objective of this study was to detect the [N II] 205 $\mu$m line, and to compare it with the [C II]157 line which has $n_{\text{rot}} = 46$ cm$^{-3}$, $T_T = 8000$ K. This has a nearly identical critical density for excitation in ionised regions. Their line ratio is directly related to the $N^+\text{C}^+$ abundance ratio, and this ratio traces the fraction of the observed [C II] emission that arises from ionized regions (Obert et al. 2006). Taking the SPIRE upper limit of 7.5 $\times$ 10$^{-8}$ W m$^{-2}$ sr$^{-1}$ with Jakob et al. (2007), the ratio of the 122/205 $\mu$m lines is $\gtrsim 1.9$, which is only adequate to constrain the ionised gas density to be $\gtrsim 30$ cm$^{-3}$. The [C II]/[N II] 205 ratio using the Jakob et al. (2007) tabulation is $\gtrsim 5.6$. Given current uncertainties and lack of an [N II]122 flux, it is necessary to await improved data. We note that the [N II] extension to the east (see Fig. 2) coincides with a hole in the excited H$_2$ emission image (Cruz-González et al. 2010), which may indicate there is a cavity of ionised gas. However clarification will require future observations with better sampling.

In Fig. 6 we show a section of the spectrum with the 398.6 $\mu$m para-H$_2$O line, and the HCO$^+$ $J = 6$–5, 7–6 and 8–7 lines. Putting the SPIRE sensitivity into perspective, Jakob et al. (2007) report that the integrated CI $^3P_1–^3P_0$ intensity measured from the KOSMA telescope with an 80$''$ beam is 46.6 K km s$^{-1}$, and main beam brightness temperature $\sim 25$ K. By comparing the same line observed with SPIRE has a peak $S/N$ ratio of $\gtrsim 14$ as seen in a single SPIRE channel. We also used RADEX to compute an LVG solution for the 2$J_1$–2$J_0$ para-H$_2$O line at 398.5 $\mu$m. Assuming similar excitation to that from the CO solution, for an abundance $X(\text{H}_2\text{O}) = 4 \times 10^{-8}$ and line width of 40 km s$^{-1}$ (Hjalmarson et al. 2003), we predict that the SPIRE flux should be $2.4 \times 10^{-8}$ W m$^{-2}$ sr$^{-1}$, which agrees with the measured value of $2.33 \pm 0.3 \times 10^{-8}$ W m$^{-2}$ sr$^{-1}$.

5. Conclusions

We have presented the SPIRE spectrum of a star-forming molecular core, DR21, showing for the first time the complete CO and $^{13}\text{CO}$ band head from $J = 4$–3 to 13–12, along with their maps at far-infrared wavelengths. A rotational temperature analysis shows two gas phases with $T_{\text{rot}}$ $\sim 80$ K and CO column density $\sim 4.5 \times 10^{18}$ cm$^{-2}$, and $T_{\text{rot}} = 185$ K and $N(\text{CO}) \sim 10^{18}$ cm$^{-2}$.

Fig. 6. Continuum subtracted spectrum showing the 398.6 $\mu$m para-H$_2$O, and HCO$^+$ $J = 7$–6 and $J = 8$–7 lines. Some instrumental features remain – notably a broad bump close to the CH$^+$ line at 347.9 $\mu$m. Although it is possible to flatten this spectrum further, we chose only to fit a third order polynomial to the entire 671–319 $\mu$m range of the SLW detector.

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References


